

LETTER TO THE EDITOR

Detection of barium in the atmospheres of the ultra-hot gas giants WASP-76b and WASP-121b*

Together with new detections of Co and Sr+ on WASP-121b

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ABSTRACT

Context. High-resolution spectroscopy studies of ultra-hot Jupiters have been key in our understanding of exoplanet atmospheres. Observing into the atmospheres of these giant planets allows for direct constraints on their atmospheric compositions and dynamics while laying the groundwork for new research regarding their formation and evolution environments.

Aims. Two of the most well-studied ultra-hot Jupiters are WASP-76b and WASP-121b, with multiple detected chemical species and strong signatures of their atmospheric dynamics. We take a new look at these two exceptional ultra-hot Jupiters by reanalyzing the transit observations taken with ESPRESSO at the Very Large Telescope and attempt to detect additional species.

Methods. To extract the planetary spectra of the two targets, we corrected for the telluric absorption and removed the stellar spectrum contributions. We then exploited new synthetic templates that were specifically designed for ultra-hot Jupiters in combination with the cross-correlation technique to unveil species that remained undetected by previous analyses.

Results. We add a novel detection of Ba+ to the known atmospheric compositions of WASP-76b and WASP-121b, the heaviest species detected to date in any exoplanetary atmosphere, with additional new detections of Co and Sr+ and a tentative detection of Ti+ for WASP-121b. We also confirm the presence of Ca+, Cr, Fe, H, Li, Mg, Mn, Na, and V on both WASP-76b and WASP-121b, with the addition of Ca, Fe+, and Ni for the latter. Finally, we also confirm the clear asymmetric absorption feature of Ca+ on WASP-121b, with an excess absorption at the bluer wavelengths and an effective planet radius beyond the Roche lobe. This indicates that the signal may arise from the escape of planetary atmosphere.

Key words. Planets and satellites: atmospheres – Planets and satellites:composition – Planets and satellites:gaseous planets – Techniques: spectroscopic – Planets and satellites: individual: WASP-76b, WASP-121b

1. Introduction

Ultra-hot Jupiters are currently the most readily accessible laboratories for the study of exoplanet atmospheres. Their size and large atmospheric scale heights, combined with the proximity to the host stars, make them appealing targets for the study of light that is transmitted through planetary atmospheres. With the recent developments of instruments for high-resolution spectroscopy (e.g., ESPRESSO; Pepe et al. 2010, 2021), the enhanced ability to retrieve high-resolution planetary spectra from the transit observations of ultra-hot Jupiters provided unique glimpses into the atmosphere of these extreme worlds. From the detection of several chemical species (e.g., Casasayas-Barris et al. 2019; Tabernero et al. 2021) to evaporating atmospheres (e.g., Yan & Henning 2018) and the study of winds (e.g., Ehrenreich et al. 2020; Seidel et al. 2021), resolving line features over

short exposures has proven to be key to unraveling these distant alien atmospheres.

Two of the best examples are WASP-76b and WASP-121b (West et al. 2016; Delrez et al. 2016). Both planets are inflated ultra-hot Jupiters on orbits with periods shorter than two days and equilibrium temperatures close to 2500 K (for more information, see Appendix B.1). They were observed with ESPRESSO, and on one of the observing nights, the four UTs of the Very Large Telescope (VLT) were used. These targets are currently benchmarks in the study of atmospheric composition and dynamics (Seidel et al. 2019; Ehrenreich et al. 2020; Tabernero et al. 2021; Fortney et al. 2021; Sánchez-López et al. 2022). Transmission spectroscopy studies report the detection of species from H¹, Li, Na, Mg, K, Ca+, V, Cr, Mn, Fe, Co¹, Ni, and Sr+ for WASP-76b (Tabernero et al. 2021; Kesseli et al. 2022) and H, Li, Na, Mg, K, Ca, Ca+, Sc+, V, Cr, Mn¹, Fe, Fe+, and Ni for WASP-121b (Hoeijmakers et al. 2020; Borsa et al. 2021; Merritt et al. 2021) using ESPRESSO high-resolution observa-

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¹ For these species, only tentative detections are given. We confirm H in WASP-76b and Mn in WASP-121b as detected.

Table 1: Summary of the transit observations of WASP-76b and WASP-121b.

Target	Night	Exp. time (s)	N _{obs}	N _{obs, in-transit} ^a	Airmass change ^b	S/N@550nm	Mode
WASP-76b	2018-09-03 (N1)	600	35	16 (20)	1.99 - 1.13 - 1.37	~101	1-UT HR21
WASP-76b	2018-10-31 (N2)	300	70	30 (37)	1.42 - 1.13 - 2.64	~79	1-UT HR21
WASP-121b	2018-11-30 (N1)	300	29	17 (20)	1.11 - 1.03 - 1.07	~151	4-UT MR42
WASP-121b	2019-01-06 (N2)	400	52	17 (23)	1.20 - 1.03 - 1.61	~47	1-UT HR21

Notes. ^a Number of spectra obtained between second and third contact. In parentheses, we list the number for the duration of the first to fourth contact. ^b Airmass change during the night (start - minimum - end).

tions for two transits each. The detection of these species, together with their respective velocities, broadening, and depths, provides important insights into the composition and dynamics of these extreme atmospheres.

In this study we revisit these datasets and extend the list of currently detected species. This paper is structured as follows. In Section 2 we describe the observations, and in Section 3 we show how we reduced the data, extracted the planetary spectra, and cross-correlated these spectra with synthetic spectra that were specifically designed for ultra-hot Jupiters. The results and discussion are then given in Section 4.

2. Observational data

We analyzed data from two transits of WASP-76b and WASP-121b that were observed with ESPRESSO. These same datasets have recently been studied and led to the detection of multiple species for each of the planets (WASP-76b: Tabernero et al. 2021; Kesseli et al. 2022; WASP-121b: Borsa et al. 2021). The 1UT observations we used were obtained as part of the ESPRESSO Guaranteed Time Observations (program 1102.C-0744, PI: F. Pepe) that cover a wavelength range from 3800 Å to 7880 Å. The HR21 mode (high resolution with 2×1 binning, Pepe et al. 2021) was used, and a resolution of $R \sim 140\,000$ was achieved. The 4UT observations were taken during the commissioning of ESPRESSO and used the MR42 mode (medium resolution with $4 \times$ and a $2 \times$ binning in the spatial and dispersion directions, $R \sim 70\,000$). The observations are summarized in Table 1. For more information, we refer to Tabernero et al. (2021) and Borsa et al. (2021).

The data were reduced using version 2.2.8 of the Reduction Software (DRS) pipeline². From the obtained data products, we used the sky-subtracted 1D (orders merged) spectra S1D_SKYSUB_A.

3. Planetary transmission spectrum and cross-correlation analysis

3.1. Telluric correction

Before light reaches the VLT mirrors, it crosses Earth’s atmosphere, which leaves its imprints on the observed spectra in the form of terrestrial absorption features. To correct for Earth absorption lines, we use the Molecfit³ pipeline (Smette et al. 2015; Kausch et al. 2015) within the ExoReflux environment (Freudling et al. 2013) as in Allart et al. (2017). With this same

tool, we simultaneously accounted for the barycentric Earth radial velocity (BERV) in each exposure (Allart et al. 2017). The wavelength regions we used to fit the terrestrial features are 6890-6900 Å, 7160-7340 Å, and 7590-7770 Å (as in Tabernero et al. 2021). These regions were selected because they are rich in telluric lines and poor in stellar features. We alternatively attempted to optimize these regions by defining more smaller regions and were very careful to exclude any stellar features. However, the two reductions produced similar results (illustrated in Fig. C). In Appendix C we provide our Molecfit input parameters together with an illustrative plot of before and after telluric correction around a single line of the sodium doublet (5891.58 Å, Fig. C.1).

3.2. Planetary spectrum extraction

To extract the planetary transmission spectrum, we followed a reasoning similar to the techniques outlined by Redfield et al. (2008) and Wyttenbach et al. (2015). These techniques have been successfully applied and improved in many recent studies using high-resolution transmission spectroscopy (e.g., Yan et al. 2017; Allart et al. 2020; Tabernero et al. 2021; Borsa et al. 2021; Casasayas-Barris et al. 2021; Seidel et al. 2022).

We started by shifting each individual spectrum into the reference frame of the star. By applying Keplerian models, we corrected for the Doppler shifts arising from the orbital pull of the planet. We then created a template of the star without any planetary absorption. To do this, we combined the out-of-transit exposures into a stellar spectrum with a high signal-to-noise ratio, called the master-out. To build this master-out, we first corrected the shape of the retrieved continuum changes for flux variations across the exposures that are due to different atmospheric conditions (airmass, seeing, etc.). We started by selecting the spectrum with the lowest airmass as the reference, and divided each of the individual spectra by this reference. We then masked the region of the spectrum in which the signal-to-noise ratio was low⁴ or that included strong telluric features⁵ and fit a third-degree polynomial function to the unmasked data. We divided out the trend for each of the individual spectra and combined the corrected out-of-transit spectra with a weighted average to create a preliminary master-out, using the inverse square of flux uncertainties as the weights.

We then performed a similar procedure, in which we instead divided the individual spectra by the created preliminary master-out. The reason for this additional step was that all the individual spectra (in-transit and out-of-transit) were later divided by the same out-of-transit template. As a result of this division, a clear

² Available at <http://www.eso.org/sci/software/pipelines/index.html>

³ Available at <https://www.eso.org/sci/software/pipelines/skytools/molecfit>

⁴ Regions in which the signal-to-noise ratio is significantly below the baseline ($> 5\sigma$), e.g., < 3900 Å.

⁵ Regions in which the telluric correction was inadequate because of low flux from strong telluric absorption, e.g., 6850-7700 Å.

wiggle pattern becomes visible across the spectra. This is to be expected for the ESPRESSO data, as noted by Tabernero et al. (2021) and Borsa et al. (2021). In Appendix D we show the wiggles across the wavelength range for a single exposure of WASP-121b with the 4UTs. The wiggle amplitude and frequency vary across wavelengths and do not behave as perfect sinusoids. To correct for these wiggles together with residual features from the flux normalization, we fit each spectrum with splines⁶. We then divided our spectra over the fit splines. An example of this fitting is illustrated in Fig. D.1.

Finally, we shifted these individual spectra according to the orbital velocity of the planet and computed the weighted averages of the spectral regions of interest. We thus created masters of the out-of-transit and in-transit data in the rest frames of the planet and star.

3.3. Cross-correlation analysis

After the planetary transmission spectrum was retrieved for each of the exposures, we performed a cross-correlation analysis with all the templates from The Mantis Network I⁷ (Kitzmann et al. 2021). These templates allow for the use of individual lines that would otherwise go unnoticed due to the high noise in the retrieved planetary spectrum. By using the cross-correlation function (CCF) approach, these individual lines can be combined into a detectable line profile. For the WASP-76b and WASP-121b line mask, we considered an isothermal atmosphere at 2500 K⁸ because of the equilibrium temperatures of the two planets (West et al. 2016; Delrez et al. 2016) and in line with the application outlined in Kitzmann et al. (2021). The number of lines used per species in our wavelength range is listed in Appendix E.

We computed the CCFs in the planet rest frame, with the previously mentioned binary masks, using the SciPy cross-correlation tool (Virtanen et al. 2020). For an easier visualization of the absorption signal, we also produced tomography plots and computed the absorption across a 2D cross-correlation grid in the Kp - planet velocity plane (hereafter Kp plots). For examples, see Fig 1.

The tomography plots show the signature of the Rossiter-McLaughlin effect (RM). Because of current limitations and uncertainties (Casasayas-Barris et al. 2021) in the modeling of the RM and the associated center-to-limb variation (CLV), we chose not to attempt to model these effects. Instead of modeling, we masked the region that matched the stellar velocities (-15 km/s to +15 km/s in the stellar rest frame) and combined the remaining signal. In addition to masking the stellar component, this process removes part of the planetary signal, which at times can cause a diminished S/N. An illustration of this masking for the F9 mask is shown in Appendix F.

⁶ We fit the splines over spectral segments of $\sim 200 \text{ \AA}$. As a first step, we also tried to fit the wiggles in each exposure using varying sinusoids on limited wavelength intervals, but we found better results with the spline method.

⁷ We performed the CCF study with templates from the following species: Li, O, Na, Mg, Al, Si, P, S, K, Ca, Ca+, Sc, Sc+, Ti, Ti+, V, V+, Cr, Cr+, Mn, Mn+, Fe, Fe+, Co, Ni, Ni+, Cu, Zn, Ga, Ge, Rb, Sr, Sr+, Y, Y+, Zr, Zr+, Nb, Mo, Ru, Rh, Pd, Cd, In, Sn, Te, Cs, Ba, Ba+, La, La+, Ce, Ce+, Pr, Nd, Nd+, Sm, Sm+, Eu, Eu+, Gd, Gd+, Tb, Tb+, Dy, Dy+, Ho, Ho+, Er, Er+, Tm, Tm+, Yb, Lu, Lu+, Hf, Hf+, W, Re, Os, Ir, Pt, Tl, Pb, Bi, Th, Th+, U, and U+. Some other species were available but had no lines in our wavelength region of interest.

⁸ For Fe+, given the empty line list at 2500 K, we used the mask corresponding to a temperature of 3000 K.

We fit Gaussian profiles to the CCFs of the species with visible planetary absorption features. We used the lmfit Python package (Newville et al. 2016) to find the best-fit values and respective uncertainties. We did not perform a CCF for hydrogen⁹, but instead directly fit the H α line on the retrieved planetary spectrum (master-in). However, we do not expect the Ca+ absorption to follow a Gaussian profile. As mentioned in Borsa et al. (2021), the Ca+ H&K lines are significantly blueshifted and their profiles are wider and deeper than those of any other species, which suggests that it extends beyond the Roche lobe¹⁰. The authors also noticed slight asymmetries in these line profiles and pointed to planetary atmospheric escape as a possible explanation. In our analysis, we found similar results and a clear asymmetric profile for the 4UT partial transit. The individual H&K lines and the combined CCF using the Mantis mask both show extended blueshifted absorption (see Fig. G.1) beyond an underlying Gaussian profile. This indicates that calcium may be escaping the planetary atmosphere. A detailed modeling of this profile might confirm this hypothesis and allow an estimate of the atmospheric escape rate of calcium.

4. Results and discussions

In Table A.1 we list the species for which a visual inspection showed tentative absorption features, together with the best-fit values and uncertainties from the Gaussian fit parameters (amplitude, center RV, and full width at half maximum, FWHM). The CCF line profile and Kp-plots for each of these are provided in Appendix H.

For WASP-76b, we find a new chemical species in the atmosphere of this planet, Ba+. It shows a strong absorption signal on both of the observing nights. We also confirm the detections made by Tabernero et al. (2021) and Kesseli et al. (2022) of H, Li, Na, Mg, Ca+, V, Cr, Mn, and Fe. The strongest K lines lie in a region that we excluded because of high telluric contamination; therefore, K was not included in our study. We did not recover Co, Ni, and Sr+ either, which were claimed by Kesseli et al. (2022). The reason might be differences in the retrieval process.

We confirm the detections of H, Li, Na, Mg, Ca+, V, Cr, Mn, Fe, and Fe+ that were claimed by Borsa et al. (2021) and of Ca and Ni by Merritt et al. (2021) in WASP-121b. The new species we claim to be present in the atmosphere of WASP-121b are Ba+, Co, and Sr+. In addition to these, we also find tentative evidence of the presence of Ti+. It was visible when we removed the signal that matched the stellar velocities (see Fig. G.2).

Overall, we see the trend that most of the detected species are blueshifted relative to their expected line positions. This trend was also observed in previous studies (e.g., Cauley et al. 2021; Pai Asnodkar et al. 2022) and is expected to be due to the winds across the terminator, which move from the dayside to the nightside of the planet.

Another surprising finding is the detection of Ba+ in both planets. This is the heaviest detected element in exoplanetary atmospheres to date. The datasets of WASP-76b and WASP-121b represent some of the highest currently available S/N datasets. That Ba+ is detected in both of the studied planets may indicate that this heavy species can be common in the atmospheres

⁹ No Mantis masks were available for hydrogen.

¹⁰ We estimated the effective planet radius at the line center by assuming $R_{\text{eff}}^2/R_p^2 = (\delta + h)/\delta$, where δ is the transit depth and h is the transmitted line amplitude (Chen et al. 2020). For Ca+ in WASP-121b, we find effective radii of $\sim 1.7 R_p$ (night 1) and $\sim 2.1 R_p$ (night 2), in agreement with Borsa et al. (2021)

of ultra-hot Jupiters. The CCF plots, tomography plots, and Kp-plots are shown in Figs. 1 and 2 for the two planets. Close to the expected planet velocity lies a clear absorption signal that follows the planetary Kp in all the exposures. Faint absorption signals directly on the planetary spectrum at the position of the strongest lines of the Ba+ CCF mask are likewise visible (Appendix I). We also note the ionization trend of these alkaline earth metals (Ca+, Sr+, and Ba+). The presence of these heavy ionized species at high altitudes in the atmospheres of ultra-hot Jupiters may be evidence of unexpected atmospheric dynamics. It is beyond the scope of this paper to describe the mechanisms that would explain the presence of these species in the upper layers of the atmosphere. However, we hope that we encourage further atmospheric modeling with this discovery.

As studied in Borsa et al. (2021), we also note the much higher S/N and clearer absorption features for the lower-resolution 4UT partial transit (despite the lower resolution: $\sim 70\,000$). This clearly demonstrates the potential of the 16m equivalent mirror or of future high-resolution spectrographs for the Extremely Large Telescope (e.g., ANDES; Marconi et al. 2021).

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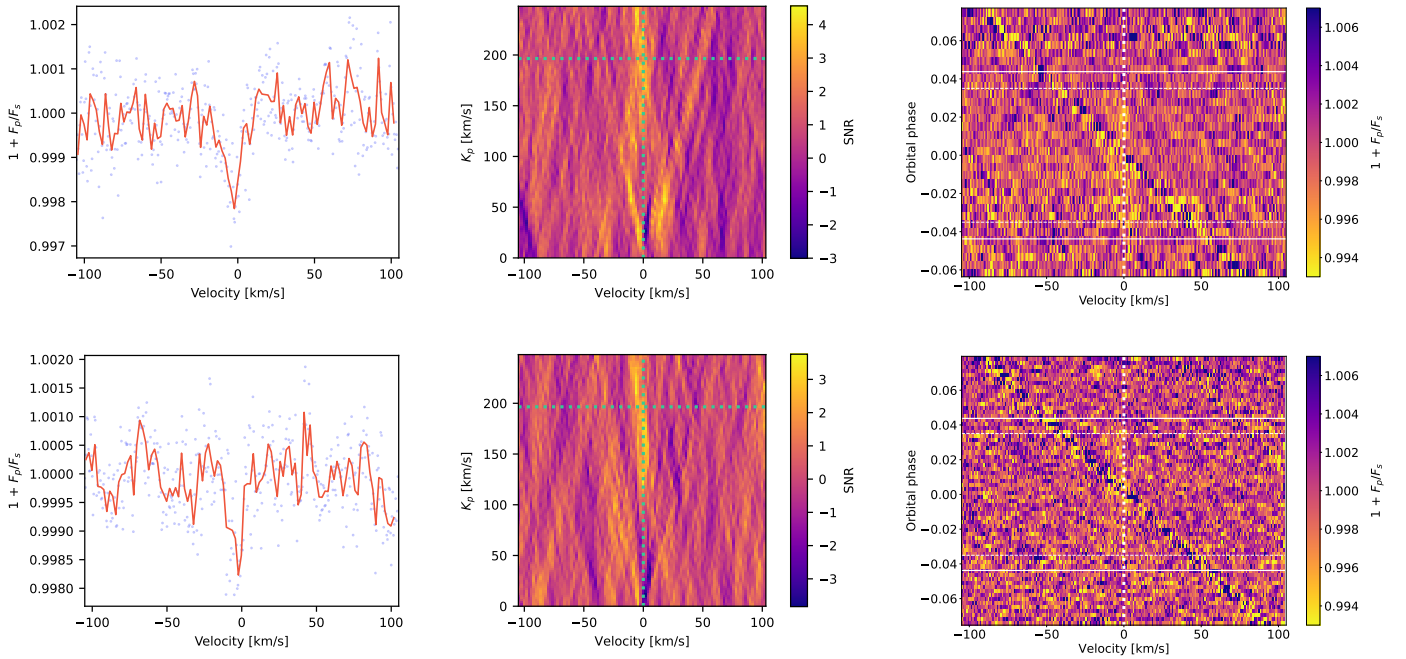


Fig. 1: Cross-correlation of the Ba+ mask with the retrieved planetary spectra for the two observation nights of WASP-76b (**top** on 2018 September 3, and **bottom** on 2018 October 31). **Left:** Cross-correlation functions of the averaged in-transit exposures on the planetary frame of reference. **Center:** Kp-plots. Map of the sum of all the individual exposures in the planet rest frame across different values in the Kp - planet velocity plane. The dashed green lines represent the expected position of the planetary signal on this map. **Right:** Tomographic plots in the planetary rest frame. The white lines represent the transit contacts.

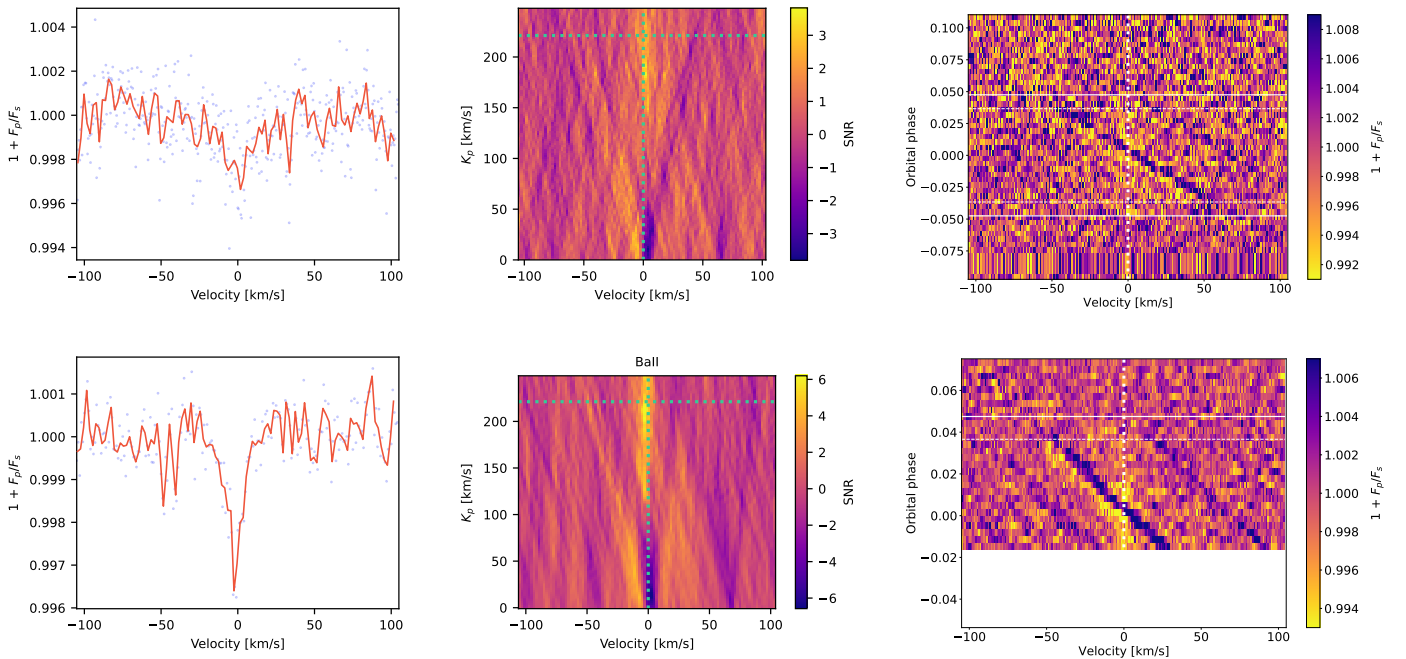


Fig. 2: Same as Fig. 1, but for the WASP-121b observing nights (**top** in 1UT mode on 2018 November 30, and **bottom** in 4UT mode on 2019 January 6).

Appendix A: Additional table

Table A.1: Summary of the Gaussian fit parameters for the detected species in the transits of WASP-76b and WASP-121b.

Species	Night	WASP-76				WASP-121			
		Amplitude [ppm]	Center RV [km s ⁻¹]	FWHM [km s ⁻¹]	S/N ^(c)	Amplitude [ppm]	Center RV [km s ⁻¹]	FWHM [km s ⁻¹]	S/N ^(c)
Ba+	N1	1803 ± 205	-3.74 ± 0.74	13.3 ± 1.8	7.0	2270 ± 296	0.5 ± 1.6	30.5 ± 5.8	6.2
	N2	1873 ± 198	-4.0 ± 0.58	10.7 ± 1.1	7.4	2926 ± 258	-1.9 ± 0.5	12.0 ± 1.3	10.6
Ca	N1	-	-	-	-	531 ± 131	0.3 ± 1.4	11.1 ± 2.9	2.6
	N2	-	-	-	-	748 ± 134	-2.13 ± 0.72	8.2 ± 1.7	6.5
Ca+ ^(a)	N1	26517 ± 1846	2.2 ± 1.4	43.1 ± 4.0	26.6	29182 ± 2620	-18.3 ± 3.0	61.1 ± 5.5	28.6
	N2	30870 ± 2527	3.0 ± 0.9	23.7 ± 2.4		54867 ± 900	-10.12 ± 0.72	92.7 ± 1.8	17.2
Co	N1	-	-	-	-	1182 ± 146	1.8 ± 0.6	9.5 ± 1.2	6.3
	N2	-	-	-	-	711 ± 88	-1.11 ± 0.94	16.7 ± 2.7	5.7
Cr	N1	383 ± 51	-5.44 ± 0.67	10.5 ± 1.7	10.0	9648 ± 72	-1.1 ± 1.0	18.3 ± 2.3	6.4
	N2	314 ± 36	-7.1 ± 1.6	30.4 ± 4.9	9.1	449 ± 60	-2.66 ± 0.88	12.2 ± 1.7	10.9
Fe	N1	667 ± 47	-3.42 ± 0.34	9.01 ± 0.68	13.5	1300 ± 72	-0.27 ± 0.32	11.1 ± 0.64	13.7
	N2	724 ± 36	-3.92 ± 0.34	13.88 ± 0.77	22.2	1141 ± 65	-2.6 ± 0.3	10.7 ± 0.68	16.3
Fe+	N1	-	-	-	-	2395 ± 366	1.1 ± 0.8	9.8 ± 1.5	4.8
	N2	-	-	-	-	2021 ± 299	-2.46 ± 0.63	8.5 ± 1.3	8.3
H ^(b)	N1	1861 ± 319	3.1 ± 4.8	43 ± 10	5.5	11085 ± 698	-2.6 ± 1.0	31.3 ± 2.0	14.8
	N2	1488 ± 373	-3.0 ± 2.7	19.2 ± 4.0	2.4	15692 ± 458	-1.57 ± 0.46	33.5 ± 1.2	46.4
Li	N1	1606 ± 153	-1.4 ± 1.2	27.6 ± 3.0	6.8	1063 ± 437	-2.9 ± 2.6	11.3 ± 4.3	2.1
	N2	1558 ± 157	-2.41 ± 0.92	17.2 ± 1.8	6.4	2163 ± 305	-8.5 ± 1.4	20.6 ± 3.2	5.2
Mg	N1	873 ± 193	-1.3 ± 2.1	17.7 ± 3.7	4.4	2821 ± 579	-2.4 ± 1.4	10.3 ± 2.8	6.2
	N2	2578 ± 237	-3.25 ± 0.57	12.2 ± 1.1	9.5	2293 ± 316	-2.37 ± 0.82	11.4 ± 1.5	6.9
Mn	N1	1307 ± 119	-3.82 ± 0.57	12.4 ± 1.2	5.9	1005 ± 242	-1.4 ± 1.0	8.0 ± 1.9	2.5
	N2	1406 ± 163	-3.94 ± 0.39	7.1 ± 1.0	7.1	1169 ± 135	-2.0 ± 0.8	13.3 ± 1.5	6.0
Na	N1	1015 ± 88	1.59 ± 0.95	20.7 ± 1.9	10.4	1406 ± 184	-0.98 ± 0.97	14.8 ± 2.2	7.8
	N2	884 ± 79	-0.0 ± 1.0	22.5 ± 2.2	8.9	1245 ± 111	-1.44 ± 0.68	15.1 ± 1.4	10.1
Ni	N1	-	-	-	-	1663 ± 522	-1.52 ± 0.24	1.94 ± 0.98	4.2
	N2	-	-	-	-	611 ± 132	-2.9 ± 1.2	11.4 ± 2.5	5.7
Sr+	N1	-	-	-	-	4548 ± 1092	1.1 ± 3.5	24.0 ± 4.4	5.7
	N2	-	-	-	-	7659 ± 781	-0.24 ± 0.85	15.6 ± 1.6	8.2
V	N1	346 ± 32	-4.88 ± 0.48	10.3 ± 1.0	6.2	618 ± 63	1.94 ± 0.65	12.8 ± 1.7	7.7
	N2	360 ± 26	-4.42 ± 0.53	14.4 ± 1.1	10.6	587 ± 41	-2.13 ± 0.44	13.0 ± 1.1	9.9

Notes. ^(a) The Ca+ signal is highly asymmetric (see Fig. G.1), therefore the Gaussian fit may not be accurate. ^(b) The H signal corresponds to the direct fit of the observed H α line in the planetary transmitted spectrum and not to a CCF, as was done with remaining species.

^(c) The S/N columns correspond to the significance of the detection. This significance was computed as the ratio of the fit amplitude to the standard deviation of the binned adjacent continuum (with the binning width equal to the fit FWHM).

Appendix B: Parameters

Table B.1: Stellar and planetary parameters.

Parameter	WASP-76		WASP-121	
	Value	Source	Value	Source
Stellar Information				
T_{eff} (K)	6329 ± 65	Ehr2020	6586 ± 59	Bou2021
$\log g$ (cgs)	4.196 ± 0.106	Ehr2020	4.47 ± 0.08	Bou2021
[Fe/H] (dex)	0.366 ± 0.053	Ehr2020	0.13 ± 0.04	Bou2021
M_* (M_{\odot})	1.458 ± 0.021	Ehr2020	1.38 ± 0.02	Bou2021
R_* (R_{\odot})	1.756 ± 0.071	Ehr2020	1.44 ± 0.03	Bou2021
Age (Gyr)	1.816 ± 0.274	Ehr2020	1.03 ± 0.43	Bou2021
V (mag)	9.52 ± 0.03	Tycho-2	10.51 ± 0.04	Tycho-2
K_s (m s^{-1})	116.02 ± 1.35	Ehr2020	0.177 ± 0.008	Bor2020
Planet Information				
Period [days]	$1.80988198 \pm 0.00000064$	Ehr2020	$1.27492504 \pm 0.00000015$	
T_0 [BJD-2400000]	$58080.626165 \pm 0.0004188$	Ehr2020	58119.72074 ± 0.00017	Bor2020
R_p/R_*	0.10852 ± 0.00096	Ehr2020	0.12534 ± 0.00005	Bor2020
a/R_*	4.08 ± 0.06	Ehr2020	3.8131 ± 0.0075	Bor2020
i [$^{\circ}$]	89.623 ± 0.034	Ehr2020	88.49 ± 0.16	Bor2020
e	0 (fixed)	Ehr2020	0 (fixed)	Bor2020
M_p (M_{Jup})	0.894 ± 0.014	Ehr2020	1.157 ± 0.070	Bor2020
R_p (R_{Jup})	1.854 ± 0.077	Ehr2020	1.753 ± 0.036	Bor2020
K_p (km s^{-1})	196.52 ± 0.94	Ehr2020	221.15 ± 17.00	Calculated

Notes. Ehr2020 - Ehrenreich et al. (2020), Bou2021 - Borsa et al. (2021), Tycho-2 - Høg et al. (2000), Bor2020 - Bourrier et al. (2020)

Appendix C: Telluric correction. Input parameters and spectrum comparison

Table C.1: Some of the key input parameters used in Molecfit.

Parameter key	Value	Description
LIST_MOLEC	H2O, O2	List of molecules to be included in the model.
WAVE_INCLUDE	0.6890,0.6900, 0.7160,0.7340, 0.7590,0.7770	Wavelength ranges to be included.
WLG_TO_MICRON	0.0001	Multiplicative factor applied to the wavelength to express it in micrometres.
WAVELENGTH_FRAME	VAC_RV	Wavelength reference frame.
FTOL	1e-8	Relative chi-square convergence criterion.
XTOL	1e-8	Relative parameter convergence criterion.
FIT_CONTINUUM	1	Flag to enable/disable the polynomial fit of the continuum.
CONTINUUM_N	4	Degree of the polynomial continuum fit to use per range.
FIT_RES_GAUSS	True	Fit resolution by Gaussian.
RES_GAUSS	3.5	Initial value for FWHM of the Gaussian in pixels, at the centre of the spectrum.
LNFL_LINE_DB	updated_aer_v_3.6 ^a	File name of the line list

Notes. Descriptions taken from the Molecfit Pipeline User Manual at <https://ftp.eso.org/pub/dfs/pipelines/instruments/molecfit/molecfit-pipeline-manual-4.1.pdf>. ^aFrom private communication with Romain Allart due to duplicated lines in the aer3.6 line list (Allart et al. 2017, 2020).

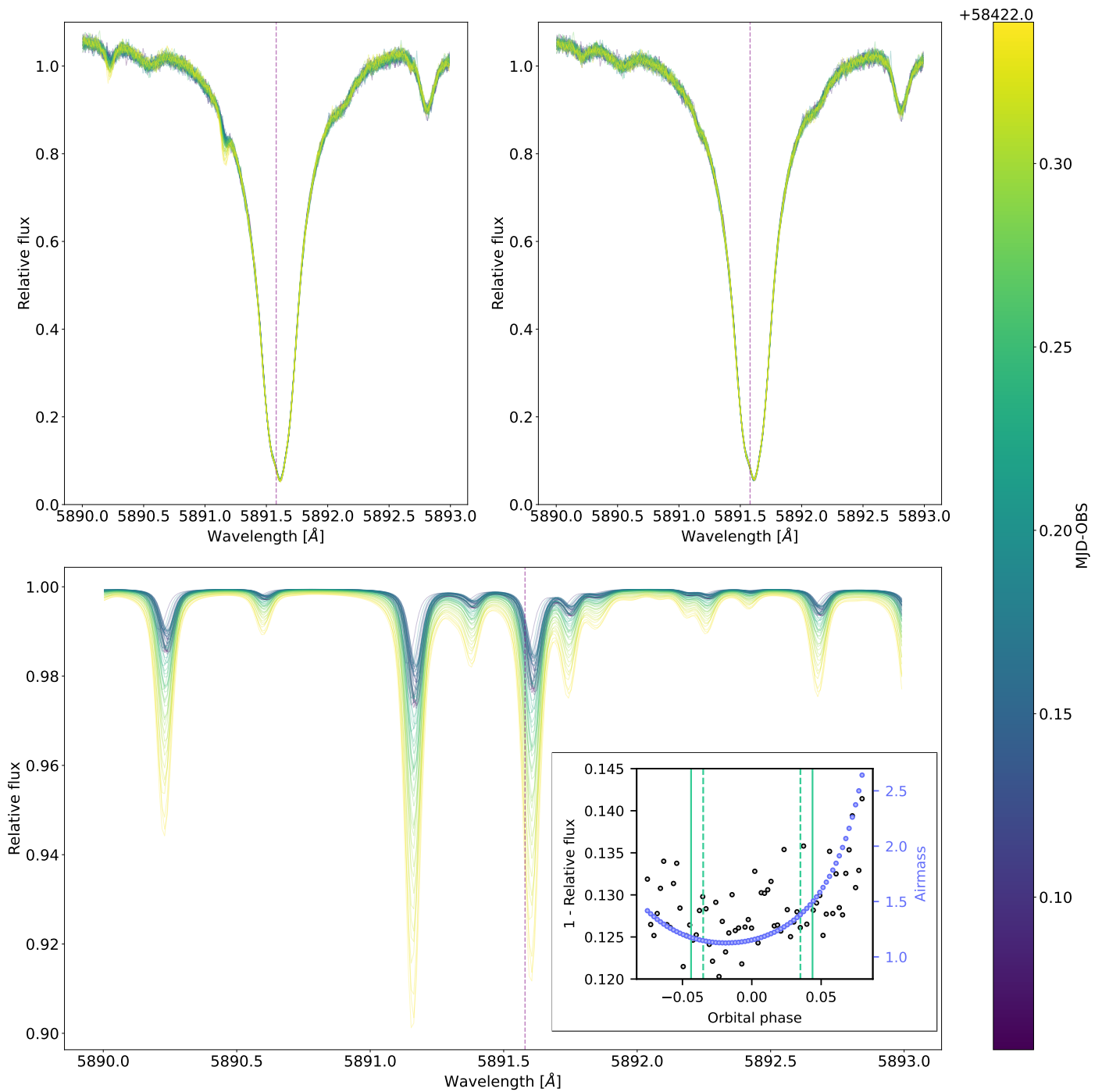


Fig. C.1: Illustration of the `Molecfit` telluric correction on a single line from the sodium D-line doublet in the observed spectrum of WASP-76b on the night of 2018 October 31 in all exposures. The vertical dashed purple line represents the expected line position. **Top left:** Sodium line prior to the telluric correction. **Top right:** Sodium line after the `Molecfit` telluric correction. **Bottom:** Ratio of the nontelluric-corrected over telluric-corrected spectral lines. **Inset:** Flux and air-mass variation over the transit observation. As expected for observations later in the night, with higher airmass, the removed telluric features are stronger.

Appendix D: Wiggles

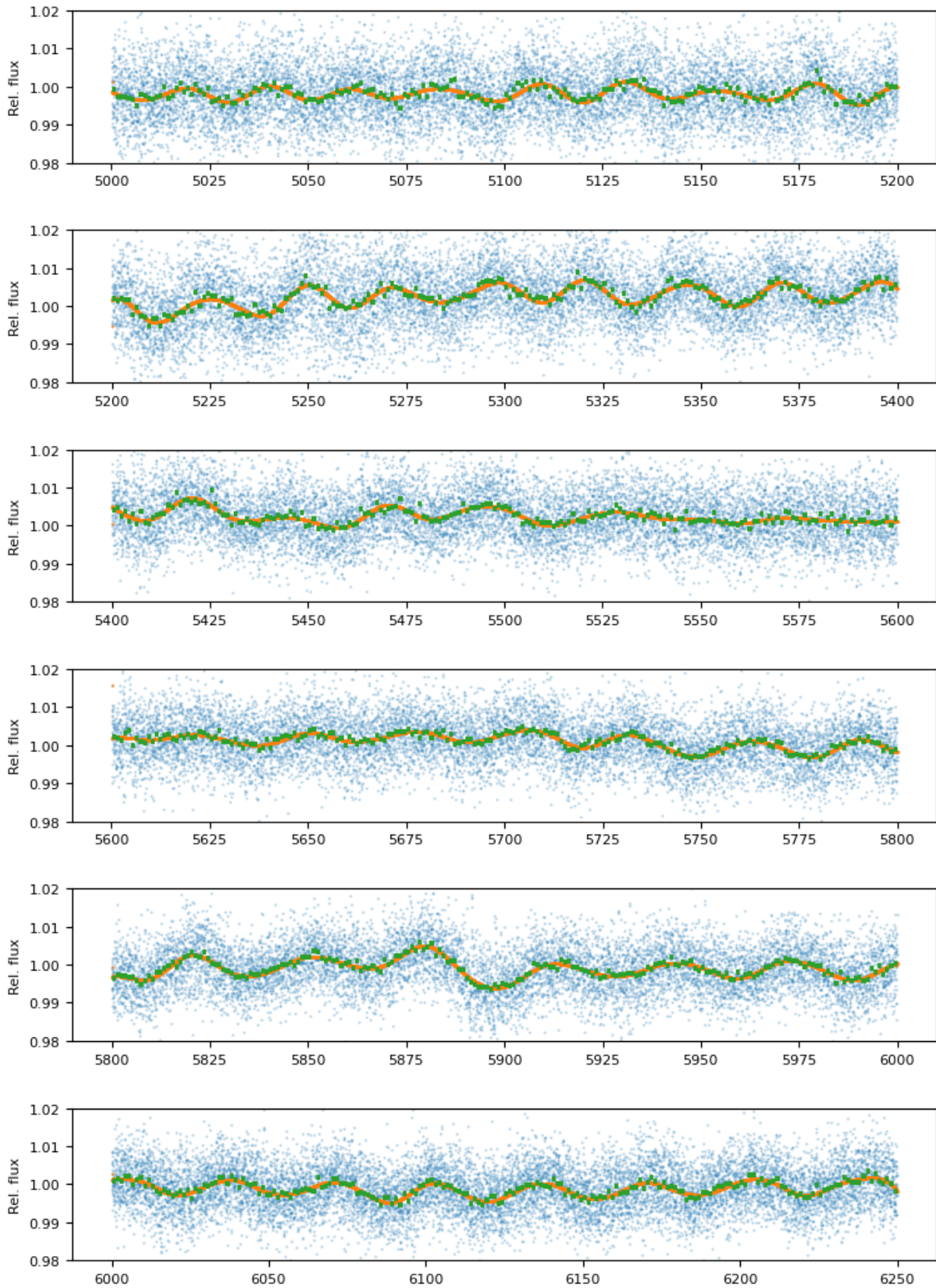


Fig. D.1: (Figure continues on the next page.)

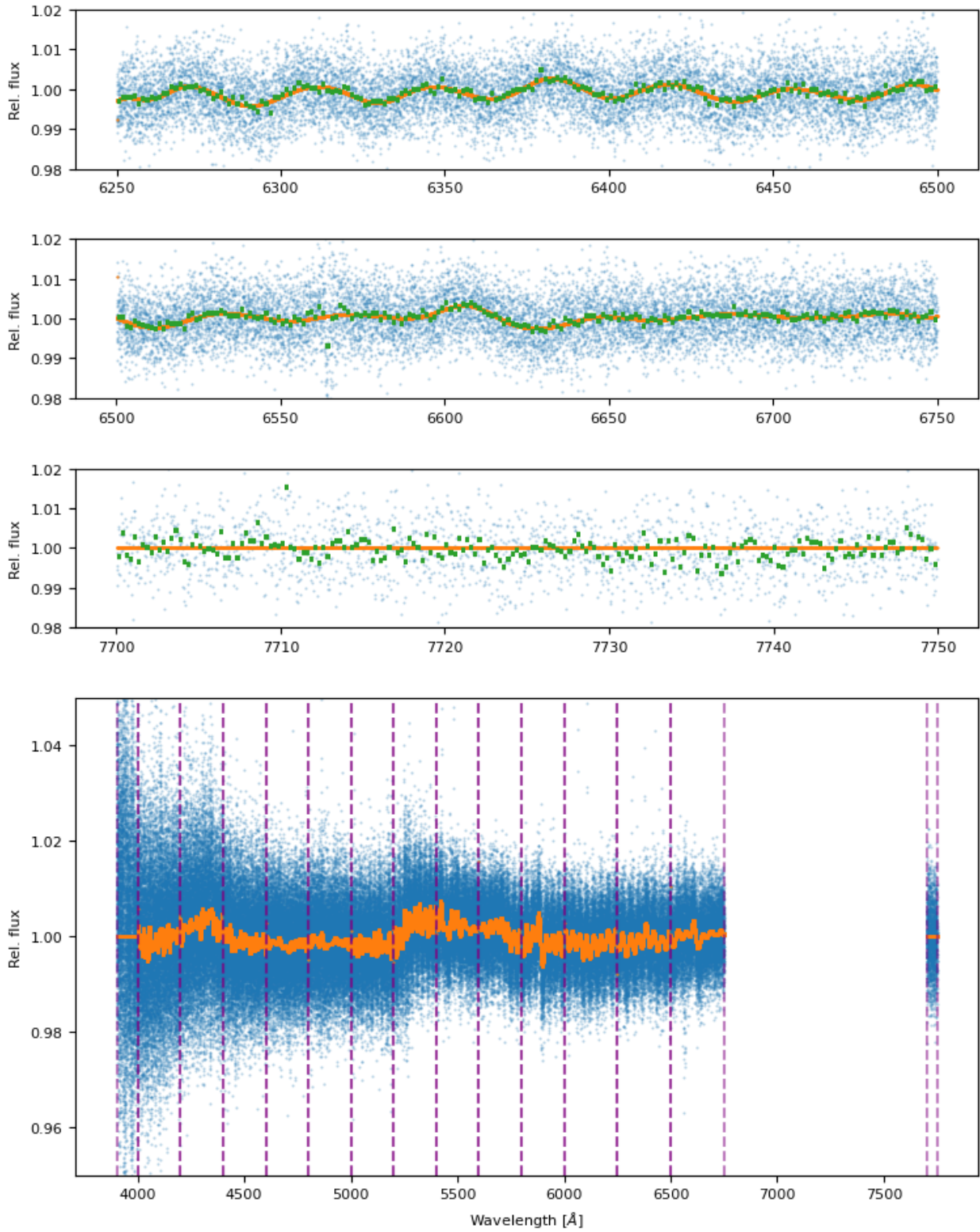


Fig. D.1: Fitting procedure during the wiggle-correction step for some of the wavelength regions in the first exposure of the 4UT observing night of WASP-121b. The fit of the entire spectral range is shown in the last panel. We divide the spectrum into regions (vertical purple lines) in which we define the priors and bounds of the fitting functions. For some regions near the lowest and highest wavelengths, we did not apply the spline fitting because it would not be a correct match to the observed data.

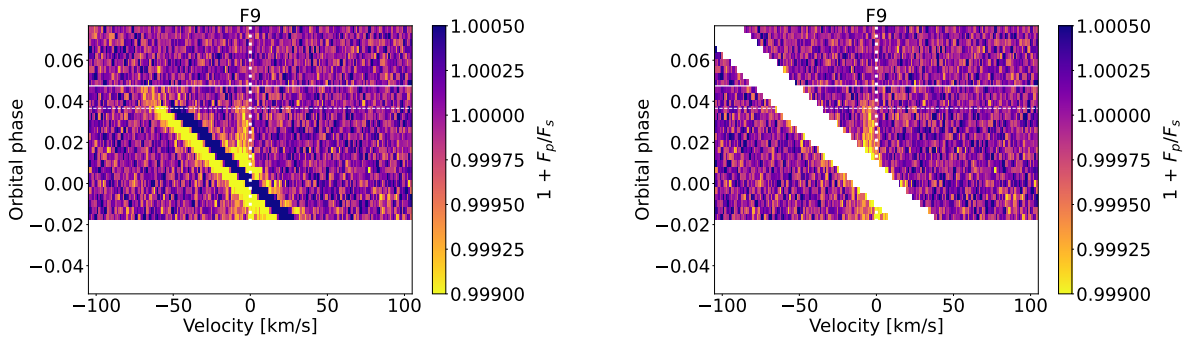


Fig. F.1: Tomography plots in the planetary rest frame built from the WASP-121b CCFs (night 2, 4UTs) of the planetary spectrum with the F9 ESPRESSO mask. **Left:** Without masking the stellar signal. **Right:** By masking the stellar velocities ranging from -15 km/s to +15 km/s in the stellar rest frame. The white horizontal band for the early orbital phases originates from the lack of data at the start of transit.

Appendix E: Lines per species used from the Mantis masks.

Table E.1: Number of lines per species we used for our CCFs using the Mantis masks at a temperature of 2500 K.

Chemical species	Number of lines
Li	9
Na	43
Mg	52
K	58
Ca	486
Ca+	2
Ti+	9
V	2543
Cr	2096
Mn	362
Fe	3827
Fe+	74 ^(a)
Co	984
Ni	580
Sr+	5
Ba+	12

Notes. ^a Lines for the 3000 K template because no lines are available at 2500 K.

Appendix F: Stellar signal masking

Appendix G: Calcium II and titanium II CCFs

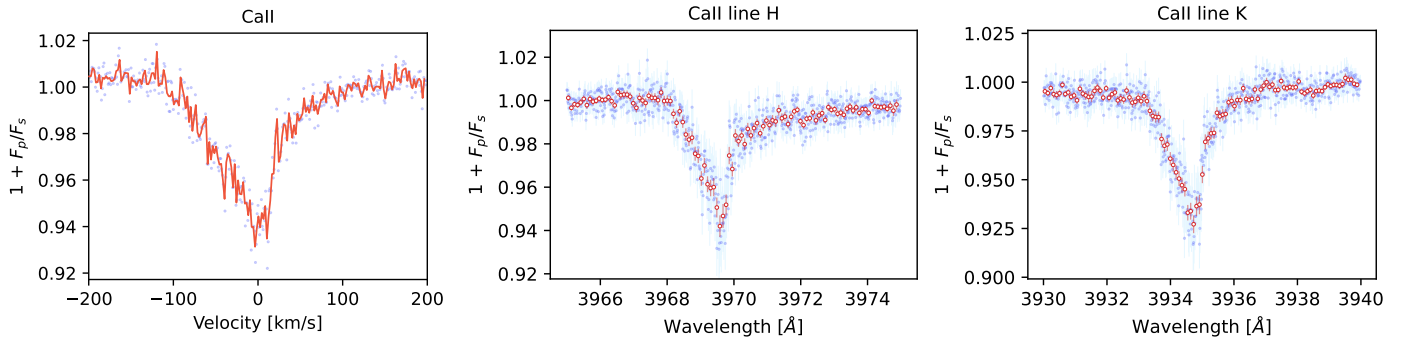


Fig. G.1: **Left:** CCF of the Ca+ mask with the retrieved planetary spectrum for WASP-121b (night 2, 4UTs) at the planetary rest frame. **Center and right:** WASP-121b retrieved planetary spectrum at the Ca+ H&K lines.

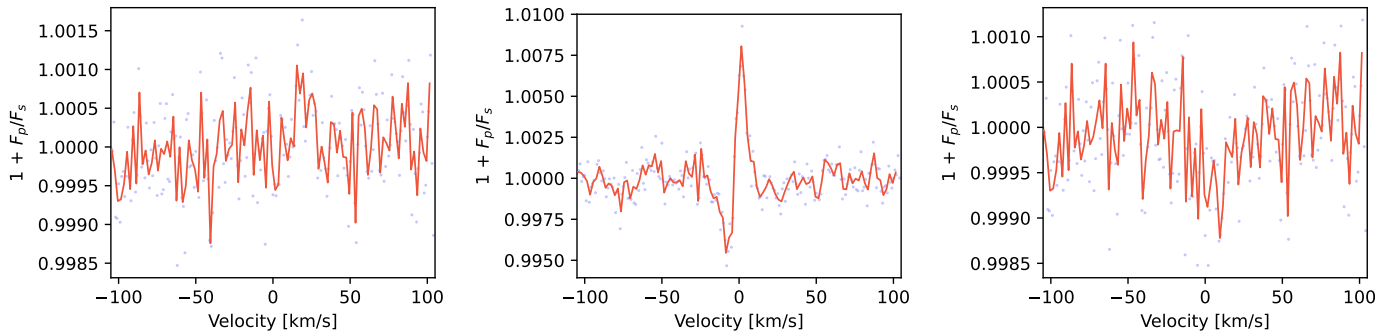


Fig. G.2: **Left:** CCF of the Ti+ mask with the retrieved planetary spectrum for WASP-121b (night 2, 4UTs). **Center:** CCF built from considering the contributions from spectra between -15 km/s and 15 km/s in the stellar frame of reference alone. **Right:** Same as the initial CCF, but excluding the previous stellar contribution region. All the plots are in the planetary rest frame.

**Appendix H: CCF and Kp-plots for WASP-76b and
WASP-121b**

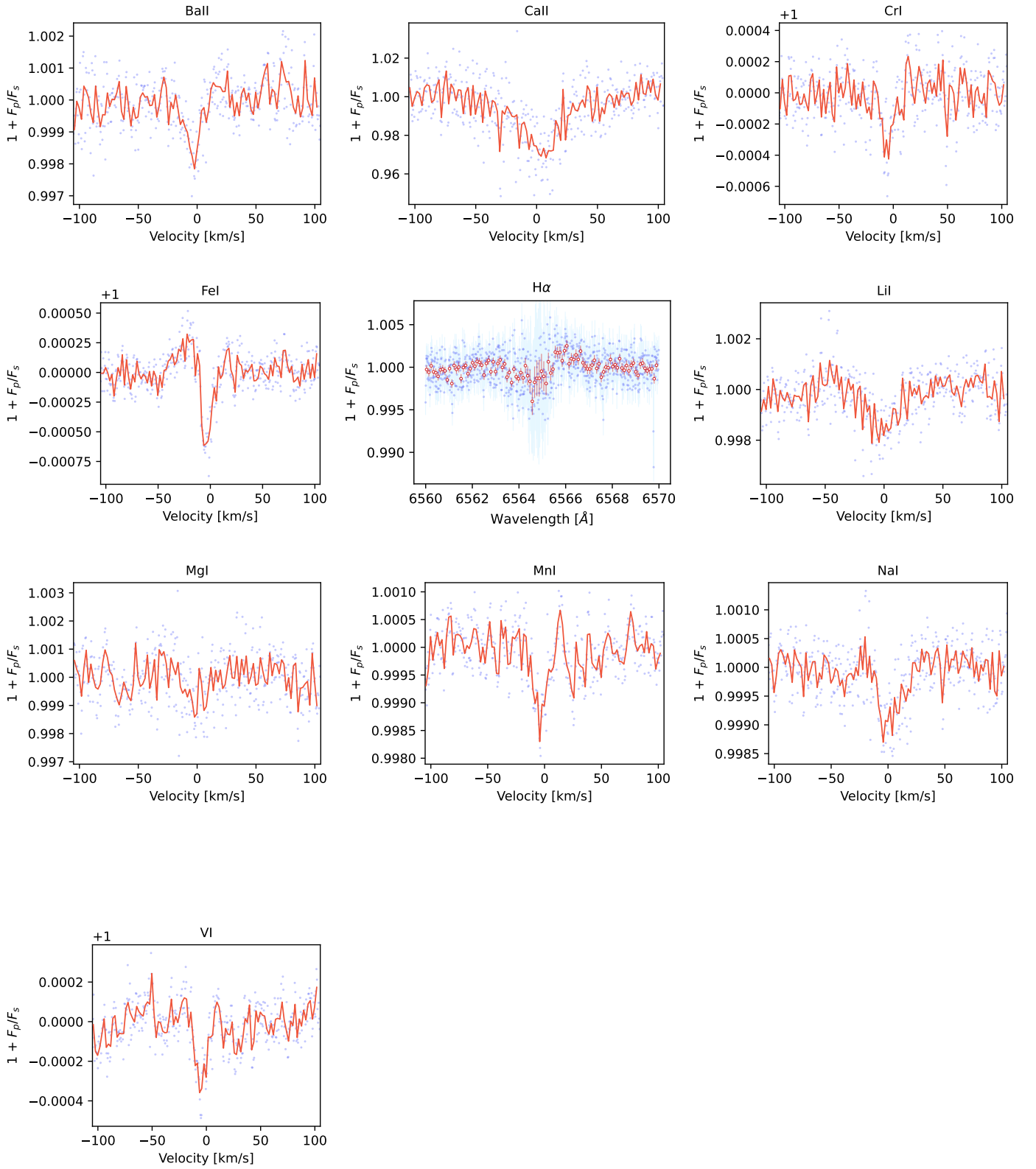


Fig. H.1: WASP-76b, night 1 (2018 September 3). CCFs of the averaged in-transit exposures in the planetary frame of reference for the detected chemical species, except for the H α line, which is present directly in the retrieved planetary spectrum.

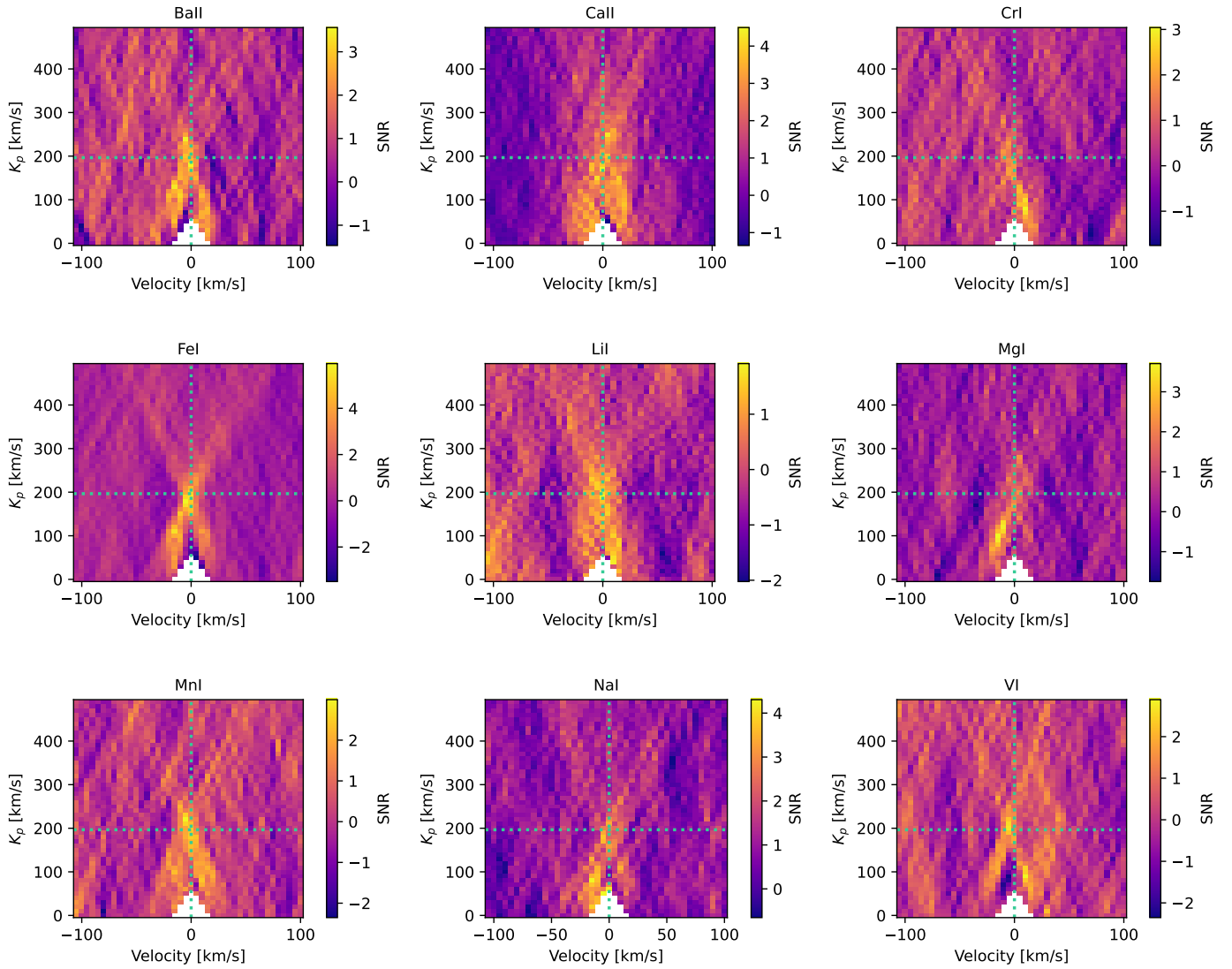


Fig. H.2: WASP-76b, night 1 (2018 September 3). Kp-plots. Map of the sum of all the individual exposures in the planetary rest frame across different values in the Kp - planetary velocity plane. We mask the signal that matches the stellar velocities ranging from -15 km/s to +15 km/s in the stellar rest frame. The dashed green lines represent the expected position of the planetary signal in this map.

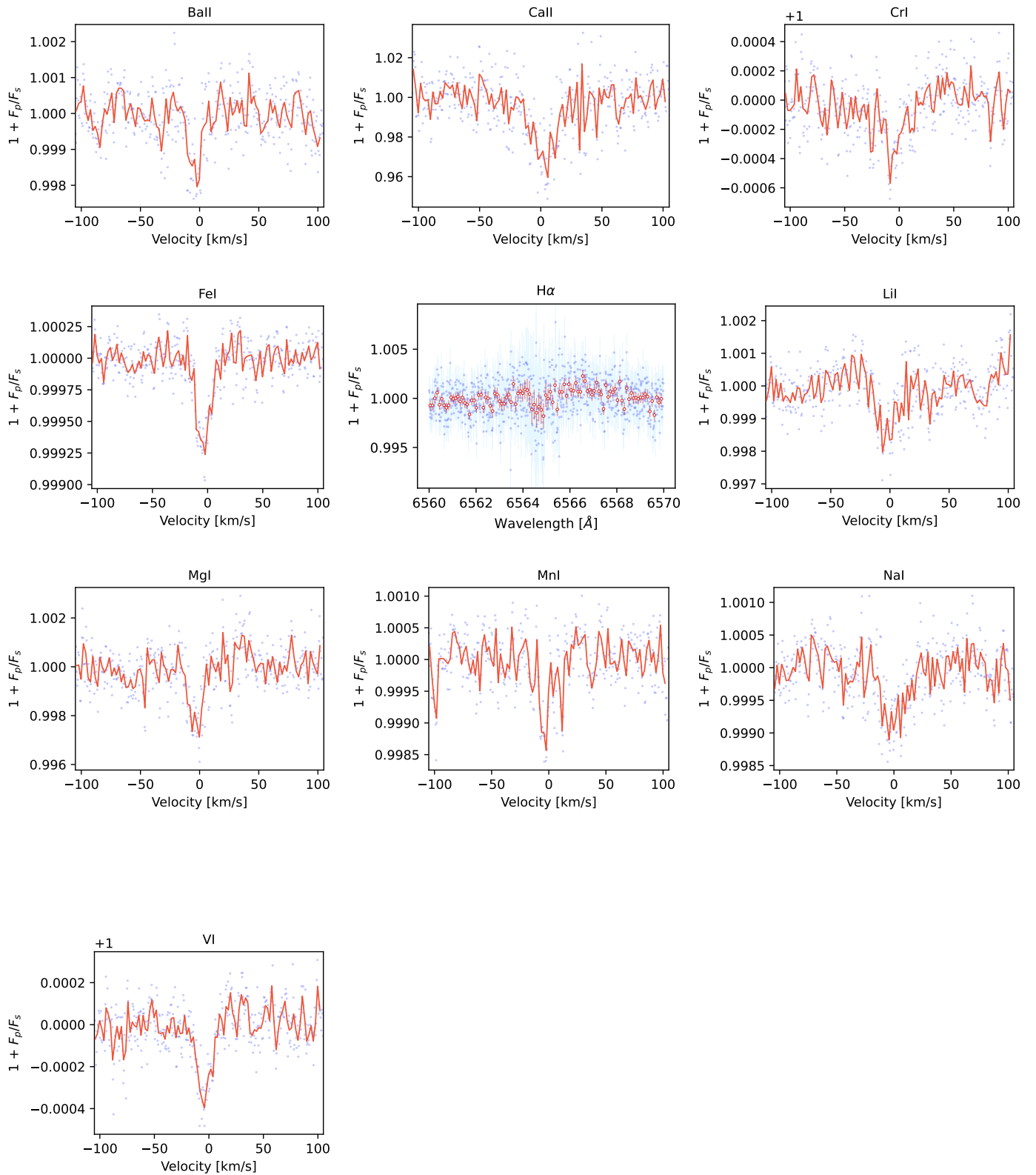


Fig. H.3: Same as Fig. H.1 for WASP-76b, night 2 (2018 October 31).

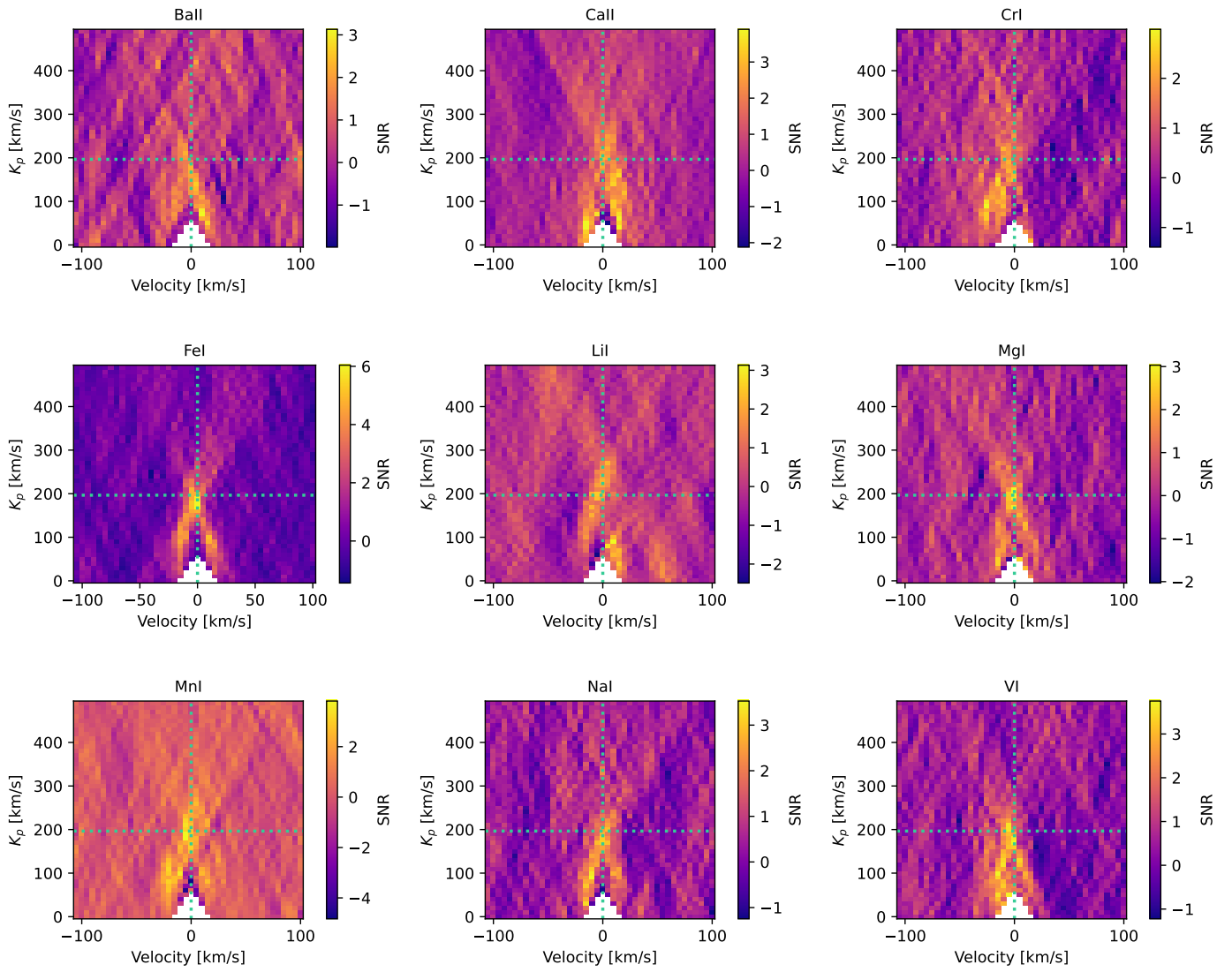


Fig. H.4: Same as Fig. H.2 for WASP-76b, night 2 (2018 October 31).

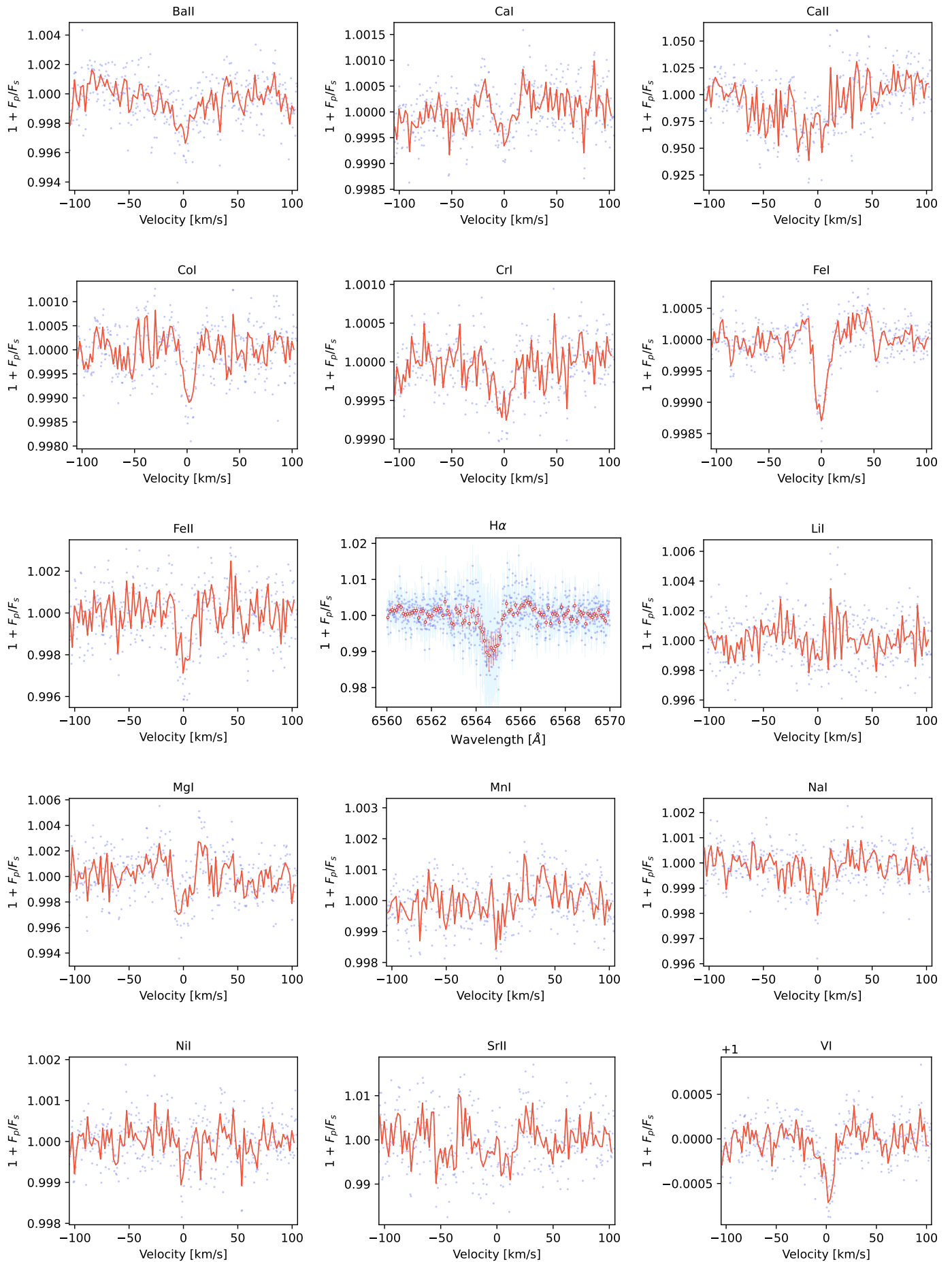


Fig. H.5: Same as Fig. H.1 for WASP-121b, night 1 (1UT - 2018 November 30).

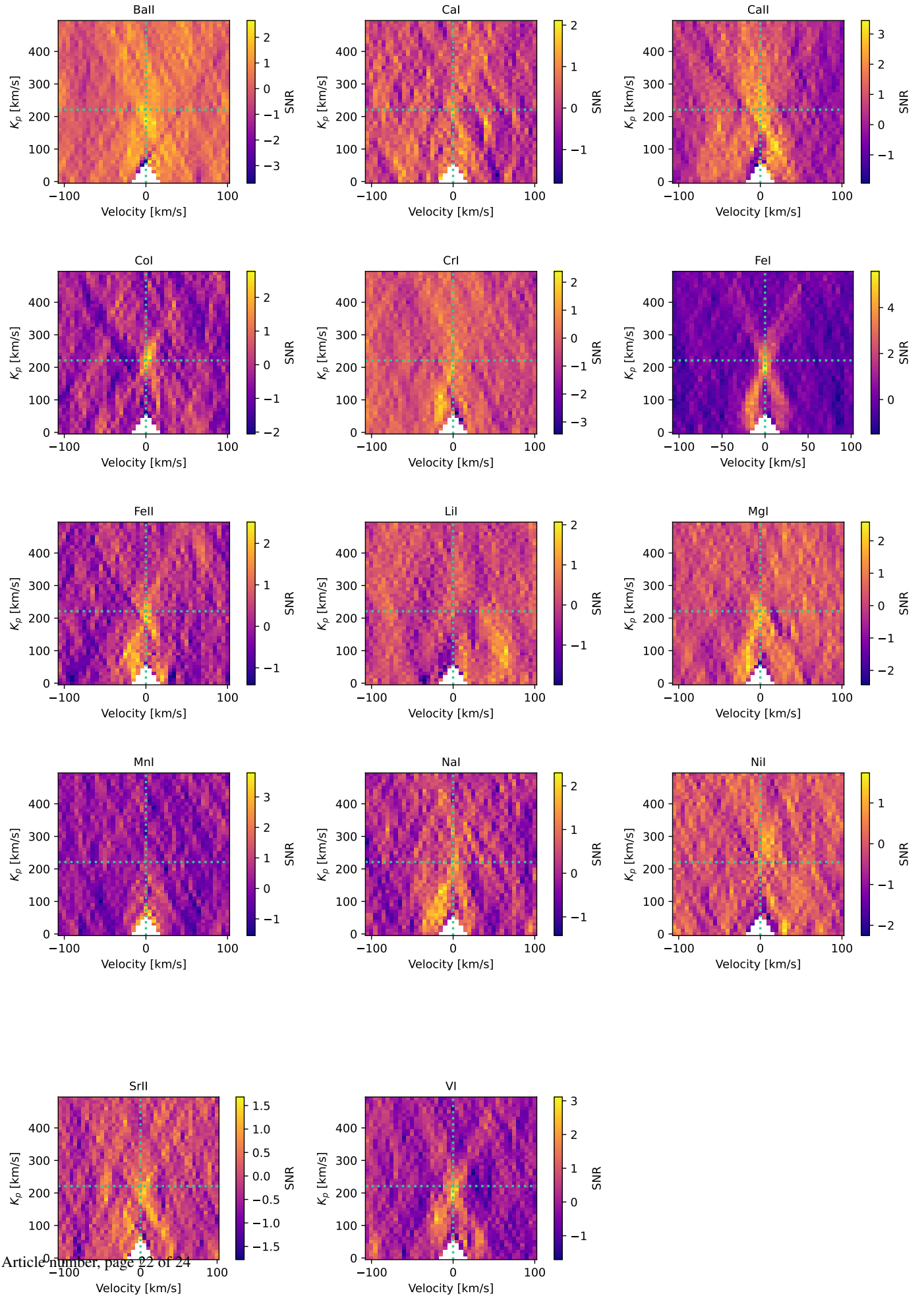


Fig. H.6: Same as Fig. H.2 for WASP-121b, night 1 (1UT - 2018 November 30).

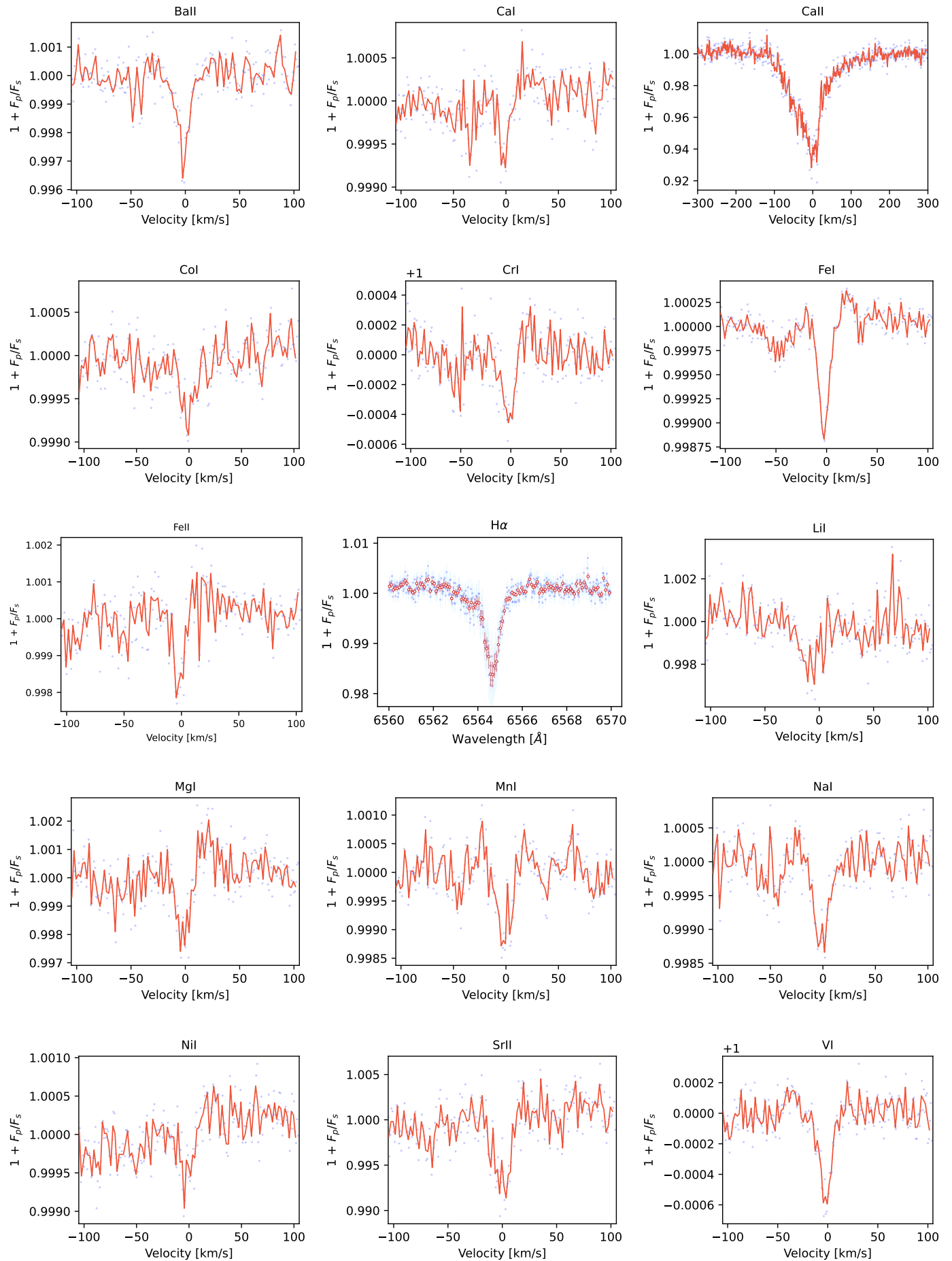


Fig. H.7: Same as Fig. H.1 for WASP-121b, night 2 (4UT's - 2019 January 6).

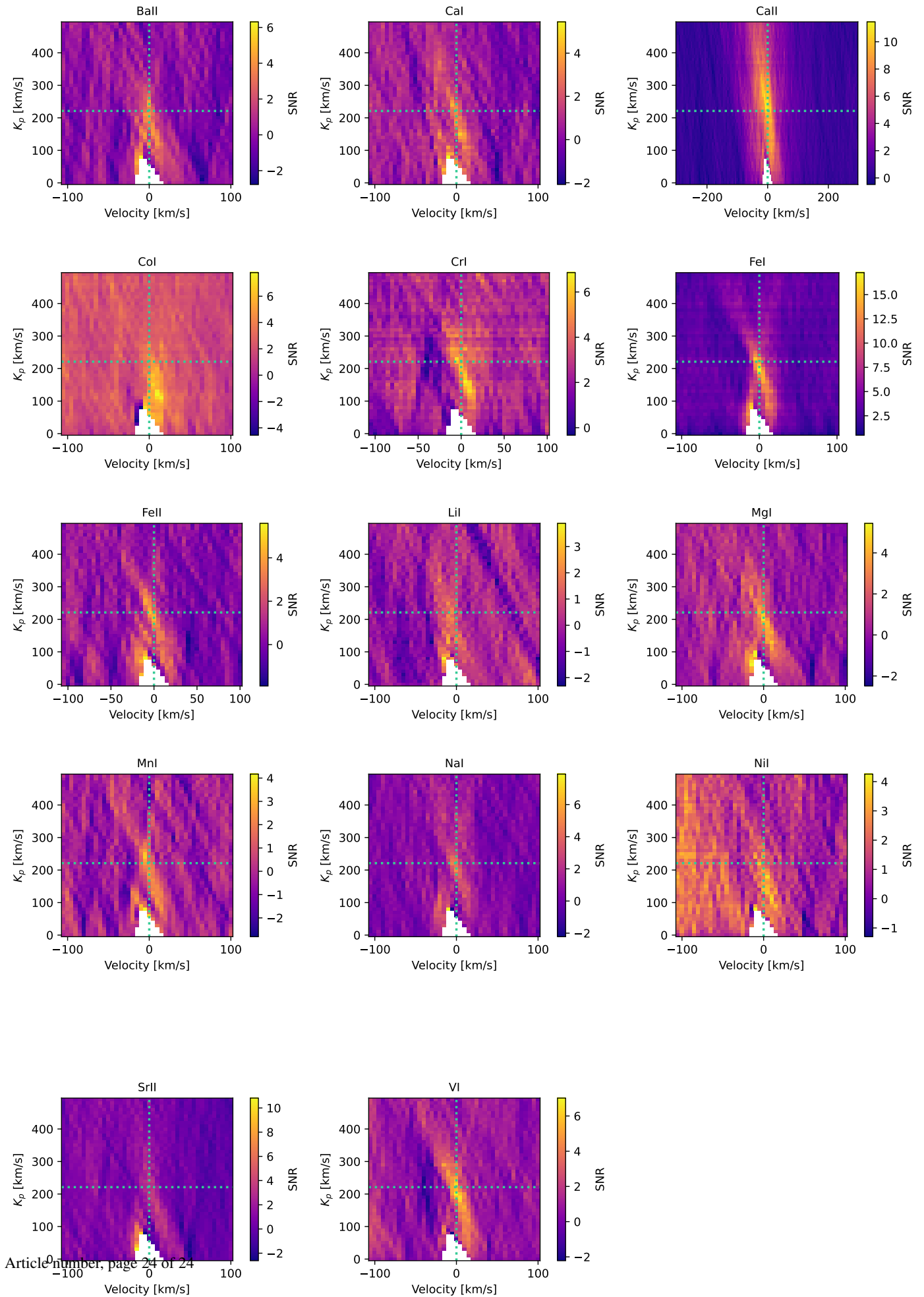


Fig. H.8: Same as Fig. H.2 for WASP-121b, night 2 (4UT's - 2019 January 6)

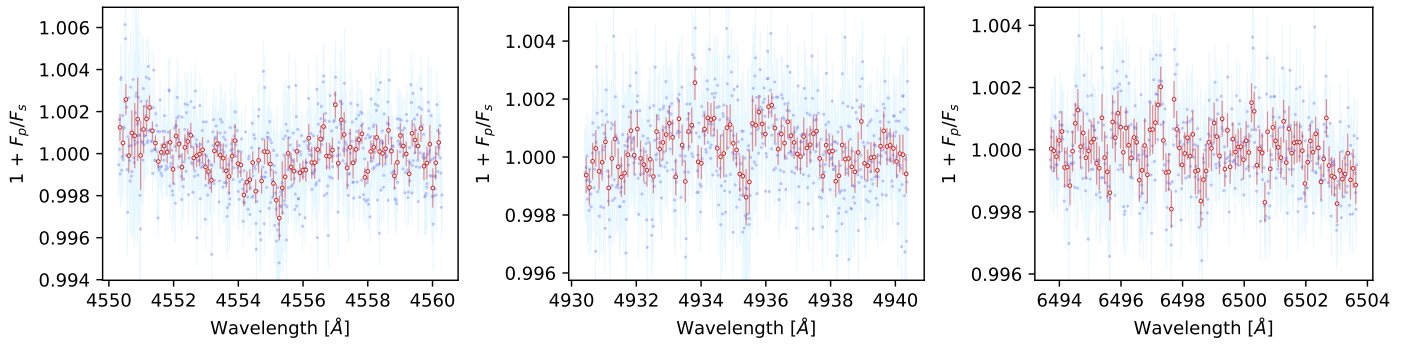


Fig. I.1: Planetary transmission spectrum for the WASP-121b, night 2 (4UTs, 2019 January 6) dataset, centered at the three strongest lines in the barium mask.

Appendix I: Individual barium II lines